

Fire, Earth, Water, Iron

Mary Bishai Brookhaven National

Neutrinos: A History

Neutrino Mixing

Main Injector Neutrino Oscillation

The Daya Bay

The Long Baseline Neutrino

Summary

Fire, Earth, Water, Iron

Harnessing the elements to study Nature's most elusive elementary particle.

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Outline

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The Long Baseline Neutrino Experiment

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- 2 Neutrino Mixing
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- 4 The Daya Bay Reactor ν Experiment
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A BRIEF HISTORY OF THE NEUTRINO



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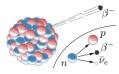
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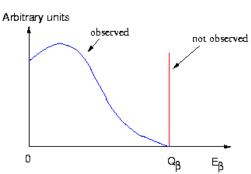
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<u>Before 1930's</u>: beta decay spectrum continuous - is this energy non-conservation?





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<u>Dec 1930:</u> Wolfgang Pauli's letter to physicists at a workshop in Tubingen:



Dear Radioactive Ladies and Gentlemen.

Wolfgang Pauli

......., I have hit upon a desparate remedy to save the "exchange theorem" of statistics and the law of conservation of energy. Namely, the possibility that there could exist in the nuclei electrically neutral particles, that I wish to call neutrons.... The mass of the neutrons should be of the same order of magnitude as the electron mass and in any event not larger than 0.01 proton masses. The continuous beta spectrum would then become understandable by the assumption that in beta decay a neutron is emitted in addition to the electron such that the sum of the energies of the neutron and the electron is constant........

Unfortunately, I cannot appear in Tubingen personally since I am indispensable here in Zurich because of a ball on the night of 6/7 December. With my best regards to you, and also to Mr Back. Your humble servant

. W. Pauli



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James Chadwick



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Solvay Conference, Bruxelles 1933: Enrico Fermi proposes to name Pauli's particle the "neutrino".



Enrico Fermi



Particle physics units and symbols

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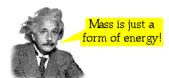
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Symbols used for some common particles:

| Symbol | Particle | | |
|------------------|-----------------------------|--|--|
| $\overline{}$ | Neutrino | | |
| γ | Photon | | |
| \mathbf{e}^{-} | Electron | | |
| e^+ | Anti-electron (positron) | | |
| р | proton | | |
| n | neutron | | |
| N | nucleon - proton or neutron | | |



Particle physicists express masses in terms of energy, mass = E/c², mass_{proton} = 1.67 \times 10⁻²⁴g \approx 1 \times 10⁹ electron-volts = 1GeV/c²

Finding Neutrinos...

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1950's: Fred Reines at Los Alamos and Clyde Cowan use the Hanford nuclear reactor (1953) and the new Savannah River nuclear reactor (1955) to find neutrinos. A detector filled with water with CdCl₂ in solution was located 11 meters from the reactor center and 12 meters underground.

The detection sequence was as follows:

$$\boxed{1} \ \bar{\nu_e} + p \rightarrow n + e^+$$

3 n +
108
 Cd \rightarrow 109 Cd* \rightarrow 109 Cd + γ ($\tau = 5\mu$ s).



Neutrinos first detected using a nuclear reactor!



ν : A Truly Elusive Particle!

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Reines and Cowan were the first to estimate the interaction strength of neutrinos.

The cross-section is $\sigma \sim 10^{-43} {\rm cm}^2$ per nucleon (p,n).

$$\nu$$
 mean free path = $\frac{\text{Mass of the proton}}{\sigma \times \text{density}}$

$$=\frac{1.67\times10^{-24}g}{10^{-43}cm^2\times11.4g/cm^3}\approx1.5\times10^{16}m=\text{1.6 LIGHT YEARS}$$

A proton has a mean free path of 10cm in lead

Neutrino detectors have to be MASSIVE



Discovery of the Muon (μ)

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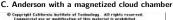
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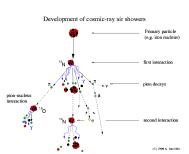
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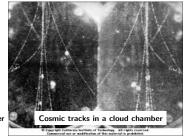
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1936: Carl Andersen, Seth Neddermeyer observed an unknown charged particle in cosmic rays with mass between that of the electron and the proton - called it the μ meson (now muons).











The Lepton Family and Flavors

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The muon and the electron are different "flavors" of the same family of elementary particles called leptons.

| Generation | | H H | III |
|-----------------|------------------|---------------------|---------------------|
| Lepton | \mathbf{e}^{-} | $oldsymbol{\mu}$ | $oldsymbol{	au}$ |
| Mass (GeV) | 0.000511 | 0.1057 | 1.78 |
| Lifetime (sec) | stable | $2.2 	imes 10^{-6}$ | 2.9×10^{-13} |

Neutrinos are neutral leptons. Do ν 's have flavor too?



Discovery of the Pion: 1947

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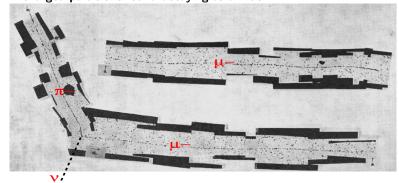
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Summary

Cecil Powell takes emulsion photos aboard high altitude RAF flights. A charged particle is found decaying to a muon:



 $\mathsf{mass}_{\pi^-} = 0.1396~\mathsf{GeV/c^2}$, $\tau = 26~\mathsf{ns}.$

Pions are composite particles from the "hadron" family which includes protons and neutrons.



Neutrinos have Flavors

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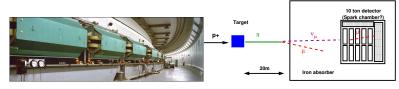
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1962: Leon Lederman, Melvin Schwartz and Jack Steinberger use BNL's Alternating Gradient Synchrotron (AGS) to produce a beam of neutrinos using the decay $\pi \to \mu \nu_x$



The AGS

Making ν 's

Result: 40 neutrino interactions recorded in the detector, 6 of the resultant particles where identified as background and 34 identified as $\mu \Rightarrow \nu_x = \nu_\mu$

The first accelerator neutrino experiment was at the AGS.



Number of Neutrino Flavors: Particle Colliders

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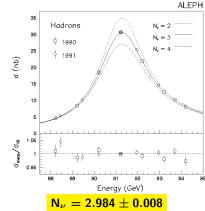
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<u>1980's - 90's:</u> The number of neutrino types is precisely determined from studies of Z^0 boson properties produced in e^+e^- colliders.

The LEP e⁺e⁻ collider at CERN, Switzerland



The 27km LEP ring was reused to build the Large Hadron Collider





The Particle Zoo

Quarks

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Forces H Higgs boson Leptons

Acknowledgements



Sources of Neutrinos

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 $\sim 1~{
m GeV}$ few/cm $^2/{
m s}$

Big Bang Reactors Sun SuperNova

10⁻⁴ eV few MeV 0.1-14 MeV ~ 10 MeV



1-20 GeV TeV-PeV $10^5/\text{cm}^2/\text{s}$ (at 1km) varies



Neutrinos and Todays Universe

Neutrino mass < 2 eV (beta-decay limits)

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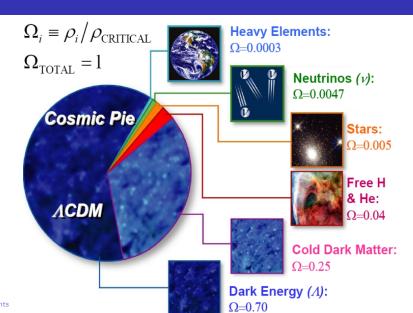
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NEUTRINO MIXING AND OSCILLATIONS



Some Quantum Mechanics

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Summary

1924: Louis-Victor-Pierre-Raymond, 7th duc de Broglie proposes in his doctoral thesis that all matter has wave-like and particle-like properties.

For highly relativistic particles : energy \approx momentum



De Broglie

Wavelength (nm)
$$\approx \frac{1.24 \times 10^{-6} \text{ GeV.nm}}{\text{Energy (GeV)}}$$

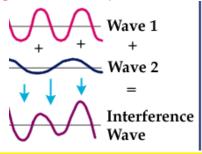


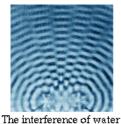
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Neutrino Mixing

1957,1967: B. Pontecorvo proposes that neutrinos of a particular flavor are a mix of quantum states with different masses that propagate with different phases:





waves coming from two sources.

The inteference pattern depends on the difference in masses



Neutrino Oscillations

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Neutrino Mixing

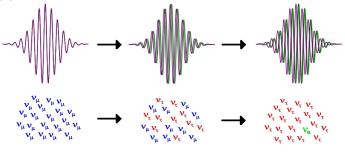
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Summary

Neutrino mass states $interfere \Rightarrow$ neutrinos oscillate between different flavors:



If neutrinos oscillate ⇒ neutrinos have mass!

The Homestake Experiment

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Summary

1967: Ray Davis from BNL installs a large detector, containing 615 tons of tetrachloroethylene (cleaning fluid), 1.6km underground in Homestake mine, SD.

1
$$\nu_{\rm e}^{\rm sun} + ^{37} {\rm CL} \rightarrow {\rm e}^- + ^{37} {\rm Ar}, \ \tau(^{37} {\rm Ar}) = 35 {\rm days}.$$

2 Number of Ar atoms \approx number of $\nu_{\rm e}^{\rm sun}$ interactions.



Ray Davis



Results: 1969 - 1993 Measured 2.5 \pm 0.2 SNU (1 SNU = 1 neutrino interaction per second for 10^{36} target atoms) while theory predicts 8 SNU. This is a $\nu_{\rm e}^{\rm sun}$ deficit of 69%.

Solar $\nu_{\rm e}$ disappearance \Rightarrow

first experimental hint of oscillations



Water Cerenkov Neutrino Detectors

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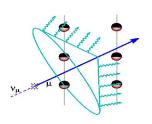
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Summary

A relativistic charged particle going through water, produces a ring of light



The Irvine-Michigan-Brookhaven Detector



With water, get MASSIVE detectors (kilo-tons)



The Super-Kamiokande Detector

Mount Ikena, Kamioka, Japan

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Neutrino Mixing

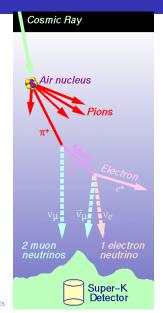
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Acknowledgements



A huge 50kT double layered tank of ultra pure water surrounded by 11,146 20" diameter photomultiplier tubes.





Identifying u_{μ} and $u_{\rm e}$ Interactions

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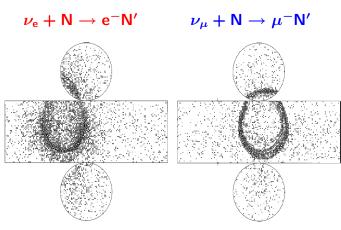


FIG. 24. An e-like (left) and a μ -like (right) event observed in the Super-Kamiokande detector.



KamLAND: Reactor $\bar{ u_e}$ Detector Kamioka Mine

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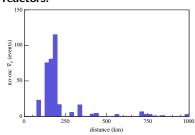
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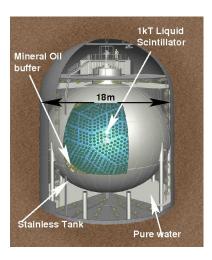
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Summai

Japan's electric power is mostly from nuclear reactors, the Kamioka Mine is bombarded by $\bar{\nu_{\rm e}}$ from reactors:







Neutrino Oscillation Results

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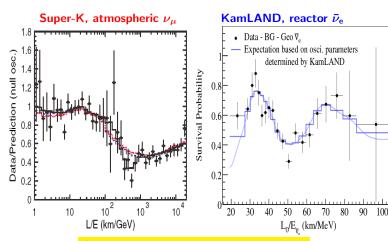
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Clear wiggles!, different beat frequencies



Neutrino Mixing: 3 flavours

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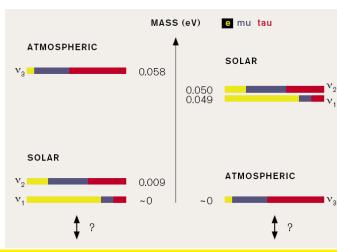
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Summary



The amount of ν_e state in ν_3 is unknown, but is constrained to be < 4%

We dont know which is larger m₃ or m₁.



Charge-Parity Symmetry

Charge-parity symmetry: laws of

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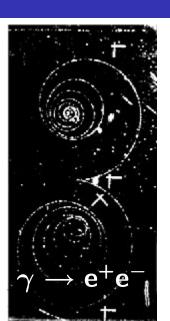
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Summary

physics are the same if a particle is interchanged with its anti-particle and left and right are swapped.

A violation of CP
matter/anti-mattery asymmetry.





Charge-Parity Symmetry and Neutrino Mixing

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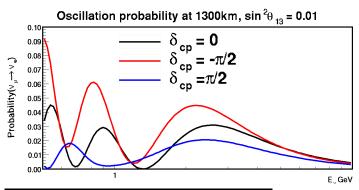
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Summary

There are 3 quantum states mixing \Rightarrow there is an overall phase: $\delta_{\rm CP}$.



If $\delta_{\mathrm{CP}}
eq 0 \ \mathrm{or} \ \pi$, charge-parity (CP) is violated

Could this explain the origin of matter?



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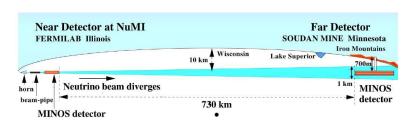
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MAIN INJECTOR NEUTRINO OSCILLATION SEARCH.





Neutrinos at the Main Injector

The longest baseline accel. ν expt in operation. Average power = 320 kW.

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Fermi Natl. Lab.. IL



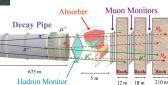
Main Injector

10 m 30 m



NuMI Horn 2 inner conductor Radial field. B ∝ 1/r

3T at 200 kA





The MINOS Collaboration

BNL: Mary Bishai, Mark Dierckxsens, Milind Diwan, David Jaffe, Brett Viren, Lisa Whitehead, Kevin Zhang

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140 scientists, 31 institutions, 6 countries



The MINOS Detectors

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Summary

Far Detector

Near Detector

Vetr shreto

electronic racks

- 484 octogonal steel and scintillator plates 8m wide,
 - \Rightarrow 5.4kTon and 30 m in length .
- Magnetized
- Cosmic µ veto shield

- 282 "squashed" octagonal steel plates, 153 scintillator planes.
 - \Rightarrow 1kTon and 16 m in length .
- Magentized



MINOS Near Detector Event Display

Brett Viren, BNL, lead developer

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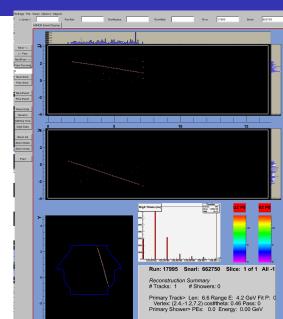
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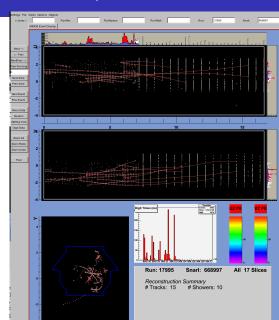
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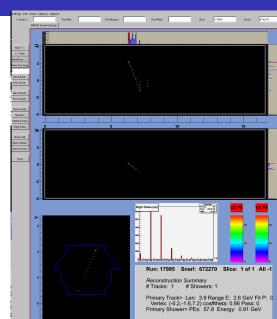
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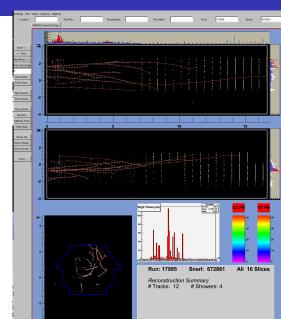
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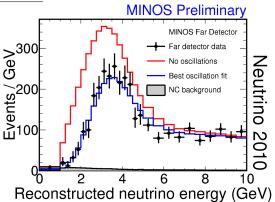
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u}$ Experiment

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Expected no-osc 2451.

Observe 1986.



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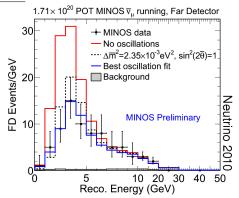
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$ar{ u}_{\mu}$ disappearance



Expected no osc 155.

Observe 97.

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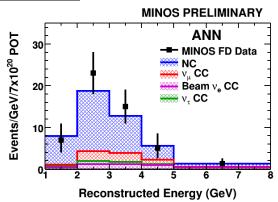
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 $\nu_{\rm e}$ appearance NEW 2010



Expected FD background: $49 \pm 7_{\rm stat} \pm 3_{\rm sys}$. Observe 54.



MINOS Beam Performance and Simulation M. Bishai, M. Dierckxsens, Brett Viren

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Main Injector Neutrino Oscillation Search

- M. Bishai lead the BNL effort to integrate the beamline data with the detector data. \Rightarrow determine the absolute neutrino rate expected in the near and far detectors per proton pulse.
- M. Bishai has been co-convener of the MINOS Beam Systematics Group for the last 2 years. Group is responsible the NuMI beamline simulation used to extrapolate the observed near detector ν event rate to the far detector.



Proton Target Radiation Damage

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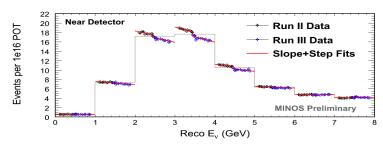
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The Daya Bay Reactor ν Experiment

The Long Baseline Neutrino Experiment

Acknowledgements

MINOS ND spectrum normalized rate vs time



Mary Bishai and David Jaffe (BNL) demonstrated that the observed drop in the ND ν rate is consistent with target damage.

Independent studies by Nick Simos (BNL-EST) of p⁺ beam irradiation damage at the BNL BLIP facility show significant damage at dosages similar to NuMI.





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Summary





THE DAYA BAY NUCLEAR REACTOR EXPERIMENT

Attempting to measure the $\nu_{\rm e}$ portion of ν_3 (AKA $\sin^2 \theta_{13}$)

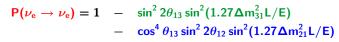


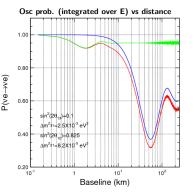
Short Baseline Reactor $\bar{\nu}_{\rm e}$ oscillations

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Acknowledgements of sin²





Unambiguous measurement

Getting to $\sin^2 \theta_{13} < 0.2\%$

Lots of statistics: -Powerful nuclear reactors + more massive detectors Supress cosmic backgrounds:

-Increase overburden = go deeper underground.

Reduce systematic uncertainties:

- -Use "identical" N/F detectors to reduce near/far detector uncertainties.
- -Calibration, calibration, calibration...



The Daya Bay Reactor Complex



Reactor Specs:

(2X2.9 GWth)

Located 55km north-east of Hong Kong.

Lina Ao II NPP (2011)

(2X2.9 GWth)

Current: 2 cores at Daya Bay site + 2 cores at Ling Ao site $= 11.6 \text{ GW}_{th}$ By 2011: 2 more cores at Ling Ao II

site = 17.4 $GW_{th} \Rightarrow top five$ worldwide

 $1~\text{GW}_{\text{th}} = 2 imes 10^{20} ar{
u_{\text{e}}}/\text{second}$

Deploy multiple near and far detectors

Reactor power uncertainties < 0.1%



The Daya Bay Collaboration

BNL: 22 collaborators from Physics and Chemistry

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Detecting Neutrinos from the Daya Bay Reactors BNL Chemistry leads scintillator effort

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Neutrino Mixing

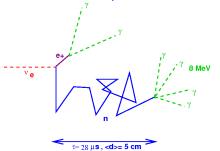
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Summary

The active target in each detector is liquid scintillator loaded with 0.1% Gd $\,$



$$\bar{\nu}_e + p \rightarrow n + e^+$$

$$ightharpoonup e^+ + e^- \rightarrow \gamma \gamma \ (2X \ 0.511 \ MeV \ +T_{e^+}, \ prompt)$$

$$\mathbf{m}$$
 n + p \rightarrow D + γ (2.2 MeV, $\tau \sim 180 \mu s$). OR

■ n + Gd
$$\rightarrow$$
 Gd* \rightarrow Gd + γ 's (8 MeV, $\tau \sim 28 \mu$ s).

 \Rightarrow delayed co-incidence of e⁺ conversion and n-capture (> 6 MeV)

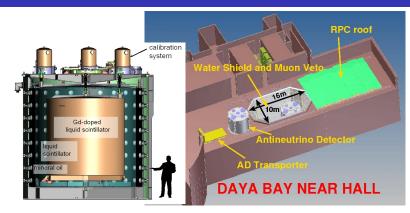


The Daya Bay Experimental Apparatus

BNL Physics lead muon veto design, and US installation/integration/safety.

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The Daya Bay Reactor ν Experiment



- Multiple "identical" detectors at each site.
- Manual and multiple automated calibration systems per detector.
- Thick water shield to reduce cosmogenic and radiation bkgds.

| | DYB | LA | Far |
|---------------------|-----|-----|-----|
| Event rates/20T/day | 840 | 740 | 90 |



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assembly pit for near/far detector pair

End 2010: Near detector ready. End 2011: far detectors ready

Reach sensitivity to $\sin^2\theta_{13} < 0.2\%$ @ 90% C.L. by 2014



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THE LONG BASELINE NEUTRINO EXPERIMENT

The search for the neutrino mass hierarchy and CP violation



Matter Effect on Neutrino Oscillation

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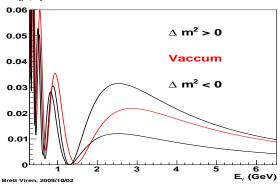
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The Long Baseline Neutrino Experiment

Summary

1978 and 1986: L. Wolfenstein, S. Mikheyev and A. Smirnov propose the scattering of ν_e on electrons in matter acts as a refrective index \Rightarrow neutrinos in matter have different effective mass than in vacuum. For $P_{\rm osc} = P(\nu_u \to \nu_e)$:

P(μ,e) at 1300 km



We can determine the mass hierarchy using $u_{\mu}
ightarrow
u_{e}$ oscillations



The Long Baseline Neutrino Experiment

247 collaborators (36 BNL), 57 institutions. Milind Diwan (BNL) co-spokesperson

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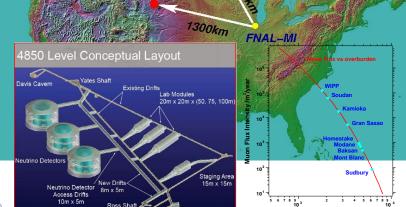
Summary

Acknowledgements

A Long Baseline Neutrino Experiment (LBNE) from Fermilab to megaton scale detectors at Homestake is now being designed. CDR late 2010.

Spuden Mine 18

Florie Stake Mine SD



Depth, meters water equivalent



Deep Underground Science and Engineering Laboratory

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Neutrinos: History

Neutrino Mixing

Neutrino Oscillation Search

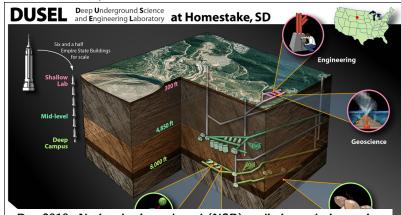
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Summary

Acknowledgements

July 10, 2007: the National Science Foundation (NSF) selected the University California-Berkeley to produce a techincal design for DUSEL at Homestake Mine, SD



Dec 2010: National science board (NSB) preliminary design review.

FY13: Earliest construction funding if approved by NSF's NSB.



Fermilab Neutrino Beams: Future

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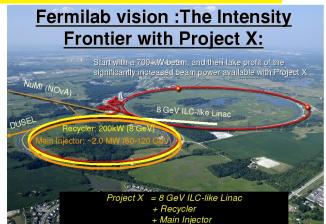
The Daya Bay Reactor ν

The Long Baseline Neutrino Experiment

Summary

The NuMI beamline uses a 300kW proton beam from the Main Injector (700 kW by 2012).

NuMI is the most powerful ν beamline operating today.



The proposed Project X at FNAL \rightarrow 2MW with E_p = 60 - 120GeV



DUSEL Detectors: Water Cerenkov BNL project management, lead detector design effort.

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Neutrinos: History

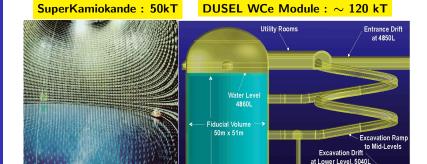
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Summary



3 100kT (fiducial) modules, \approx 55m diameter, \approx 60m height, 60K 10" PMTs/module (25% coverage)

Large Cavity

Known technology $3-4\times$ SuperK

Large backgrounds, low efficiency



DUSEL Detectors: Liquid Argon TPC BNL ITD lead effort on readout electronics

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Neutrin Mixing

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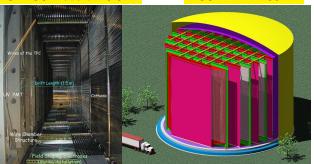
The Daya Bay Reactor ν Experiment

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Summar

ICARUS module: 0.3kT

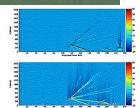
DUSEL LAr: 50 kT



ArgoNeuT (175 litre) prototype in the NuMI beam \rightarrow

High efficiency and purity

Requires $100\times$ scale-up - unproven.





LBNE/DUSEL spectra and event rates

Mary Bishai, Brett Viren, Mark Dierckxsens, BNL

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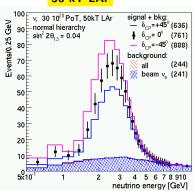
Summary

A preliminary on-axis wide-band beam for LBNE based on the NuMI focusing system has been developed. Water Cerenkov response is based on the SuperK MC. LAr is modeled as a near-perfect detector. Exposure is 3 MW. yr ν with $\sin^2 2\theta_{13} = 0.04, \; \delta_{cp} > 0, \; m_3 > m_1$

300 kT WCe

v. 30 10²⁰ PoT, 300kT WCh signal + bkg: Events/0.25 GeV normal hierarchy sin² 20, , = 0.04 (1023)background: (417)💢 beam ν_ο (202) 60 20

50 kT LAr





Summary and Conclusions

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Summary

3 flavor neutrino mixing is now well established.

BUT- what we dont know is even more important:

- What is the mass of ν_{μ} ?
- Why is the mixing in the neutrino sector so large?
- How small is $\sin^2 \theta_{13}$? Is it close to the current limit (4%) or is it very small? Is it 0?
- What is the mass hierarchy?
- Is there CP violation in the lepton sector?
- Are there only 3 generations of leptons?

In the coming decades, the Daya Bay and LBNE projects will significantly extend the neutrino discovery frontier. BNL has played a critical role in past and present ν experiments and is <u>leading</u> the charge into undiscovered territory



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Neutrinos: A

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Summary

Acknowledgeme

I'd like to thank the following members of BNL's Electron Detector Group (EDG) for their invaluable support and mentorship:

<u>Milind Diwan</u>, <u>David Jaffe</u>, <u>Brett Viren</u>, <u>Laurie Littenberg</u>,

<u>Steve Kettell</u>, Mark Dierckxsens, Lisa Whitehead.

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